

Relative magnetic helicity as a diagnostic of solar eruptivity

E. Pariat¹, J. E. Leake², G. Valori³, M. G. Linton², F. P. Zuccarello⁴, and K. Dalmasse⁵

¹ LESIA, Observatoire de Paris, PSL Research University, CNRS, Sorbonne Universités, UPMC Univ. Paris 06, Univ. Paris Diderot, Sorbonne Paris Cité, 92195 Meudon, France e-mail: etienne.pariat@obspm.fr

² Naval Research Laboratory, Washington DC, USA

³ UCL-Mullard Space Science Laboratory, Holmbury St. Mary, Dorking, Surrey, RH5 6NT, UK

⁴ Centre for mathematical Plasma Astrophysics, Department of Mathematics, KU Leuven, B-3001 Leuven, Belgium

⁵ CISL/HAO, National Center for Atmospheric Research, P.O. Box 3000, Boulder, CO 80307-3000, USA

Received ***; accepted ***

ABSTRACT

Context. The discovery of clear criteria that can deterministically describe the eruptive state of a solar active region would lead to major improvements on space weather predictions.

Aims. Using series of numerical simulations of the emergence of a magnetic flux rope in a magnetised coronal, leading either to eruptions or to stable configurations, we test several global scalar quantities for the ability to discriminate between the eruptive and the non-eruptive simulations.

Methods. From the magnetic field generated by the three-dimensional magnetohydrodynamical simulations, we compute and analyse the evolution of the magnetic flux, of the magnetic energy and its decomposition into potential and free energies, and of the relative magnetic helicity and its decomposition.

Results. Unlike the magnetic flux and magnetic energies, magnetic helicities are able to markedly distinguish the eruptive from the non-eruptive simulations. We find that the ratio of the magnetic helicity of the current-carrying magnetic field to the total relative helicity presents the highest values for the eruptive simulations, in the pre-eruptive phase only. We observe that the eruptive simulations do not possess the highest value of total magnetic helicity.

Conclusions. In the framework of our numerical study, the magnetic energies and the total relative helicity do not correspond to good eruptivity proxies. Our study highlights that the ratio of magnetic helicities diagnoses very clearly the eruptive potential of our parametric simulations. Our study shows that magnetic helicity based quantities may be very efficient for the prediction of solar eruptions.

Key words. Magnetic fields, Methods: numerical, Sun: surface magnetism, Sun: corona

1. Introduction

The reliable prediction of the triggering of solar eruptions is an essential step toward improving space weather forecasts. However, since the underlying mechanisms leading to the generation of solar eruptions have not yet been indisputably determined, no sufficient conditions of solar eruptivity have yet been established. Solar flare predictions are thus still strongly driven by empirical methods (Yu et al. 2010; Falconer et al. 2011, 2014; Barnes et al. 2016, e.g.). Nowadays most prediction relies on the determination and characterization of solar active regions through the use of several parameters and the statistical comparison with the historical eruptivity of past active regions presenting similar values of these parameters. This empirical methodology has driven the quest for the determination of which parameter, or combination of parameters, would be the best proxy of the eruptivity of solar active regions. Multiple studies have thus analysed a relatively vast list of quantities, extracting them from existing observations of active regions and comparing them with the observed eruptivity (e.g. Leka & Barnes 2003a,b, 2007; Schrijver 2007; Bobra & Couvidat 2015; Bobra & Ilonidis 2016). No simple singular parameter has yet been found to be a reliable proxy of solar eruptivity. Thanks to advances in computer science, with the development of data-mining methods and machine-learning algorithms, this search and its direct application to new flare pre-

diction systems are now reaching a new stage of development with projects such as FLARECAST¹.

The determination of pertinent parameters of eruptivity has so far been almost exclusively based on observational data. Despite tremendous advances in numerical modeling of solar eruptions, few studies have used numerical simulations to advance the search for eruptivity criteria. This is partly due to the fact that numerical models are strongly driven by either having a high level of realism (Lukin & Linton 2011; Baumann et al. 2013; Pinto et al. 2016; Carlsson et al. 2016) or focus on case-by-case studies of observed events (Jiang et al. 2012; Inoue et al. 2014; Rubio da Costa et al. 2016). Little efforts are spent on performing systematic parametric simulations which would allow the determination of eruptivity criteria. Similarly to observations, the search for proxies of eruptivity requires models which are parametrically able to produce both eruptive and non-eruptive simulations. Kusano et al. (2012) presents one of the few examples of such a simulation set-up. By varying two parameters, Kusano et al. (2012) has been able to derive an eruptivity matrix based on the relative orientation of two magnetic structures. Recently, Leake et al. (2013) and Leake et al. (2014) have presented another set of 3D MHD numerical simulations of flux emergence which, depending on a unique criterion, is able to generate either an eruptive or stable configuration.

¹ <http://flarecast.eu/>

Along with Guennou et al. (2017), the present study is dedicated to the analysis of the eruptivity properties of the simulations of Leake et al. (2013, 2014). Our goal is to determine whether a unique scalar quantity, computable at a given instant, is able to properly describe the eruptive potential of the system. While Guennou et al. (2017) will focus on the analysis of quantities than can be derived solely from the photospheric magnetic field, similarly to what is done with real observational data, the present work will be restricted to a few quantities computed from the coronal volume of the system.

Magnetic energy and magnetic helicity are typical scalar quantities that contribute to the description of a 3D magnetic system at any instant. Magnetic energy is known to be the key source of energy that fuels solar eruptions. However, only the so-called free magnetic energy (or non-potential energy), the fraction of the total magnetic energy corresponding to the current-carrying part of the magnetic field, can actually be converted to other forms of energy during a solar eruption (e.g. Emslie et al. 2012; Aschwanden et al. 2014). The non-potential energy of an active region has actually been observationally found to be positively correlated with the active region's flare index and eruptivity although not providing by itself a necessary criterion for eruptivity (Schrijver et al. 2005; Jing et al. 2009, 2010; Tziotziou et al. 2012; Yang et al. 2012; Su et al. 2014).

Magnetic helicity is a signed scalar quantity which quantifies the geometrical properties of a magnetic system in terms of twist and entanglement of the magnetic field lines. Magnetic helicity, together with magnetic energy and cross helicity, are the only three known conserved quantities of ideal MHD. Differently from the other two, however, magnetic helicity has the property of being quasi-conserved even when intense non-ideal effects are developing (Berger 1984; Parlat et al. 2015). Magnetic helicity conservation has been suggested to be the reason behind the existence of ejection of twisted magnetised structures, i.e., coronal mass ejections (CMEs), by the Sun (Rust 1994; Low 1996). There is observational evidence that magnetic helicity tends to be higher in flare-productive and CME-productive active regions compared to less productive active regions (Nindos & Andrews 2004; LaBonte et al. 2007; Park et al. 2008, 2010; Tziotziou et al. 2012). Magnetic helicity conservation can also be used to predict the outcome of different magnetic interactions, e.g., tunnel reconnection (Linton & Antiochos 2005).

The goal of the present study is thus to determine whether quantities based on either magnetic energy or magnetic helicity are able to describe the eruptivity potential of the parametric flux emergence simulations of Leake et al. (2013) and Leake et al. (2014), which we will be described in Sect. 2. In Sect. 3 will present the magnetic flux time evolution and the magnetic energy evolution in the different simulations. Sect. 4 will analyse and compare the magnetic helicity properties of the different parametric cases. In Sect. 5, we will compare the different proxies of eruptivity. We obtain the very promising results that the ratio of the current carrying magnetic helicity to the total magnetic helicity constitutes a very clear criterion describing the eruptive potential of the simulations. In Sect. 6, we will finally discuss the limitations of the present analysis and the possible application of the eruptivity criterion highlighted in this study.

2. Eruptive and non-eruptive parametric flux emergence simulations

The datasets analysed in this study are based on the 3D visco-resistive MHD simulations presented in Leake et al. (2013, hereafter L13) and Leake et al. (2014, hereafter L14). L13 and L14

have performed and analysed simulations of the emergence of a twisted magnetic flux rope in a stratified solar atmosphere. The flux tube, initially located in the high-density bottom part of the simulation box (assumed to emulate the convection zone), is made buoyant and rises up in the high-temperature corona going through a minimum temperature region emulating the solar photosphere. In all the seven simulations considered here, the initial conditions constituting the emerging flux rope and the thermodynamical properties of the atmosphere are kept strictly constant. The atmospheric stratification and the initial condition of the flux rope are typical of emergence simulations, and are presented in details in L13 and L14. The simulations are also performed with the very same numerical treatment. The employed mesh is a 3D irregular Cartesian grid with z corresponding to the vertical direction (the gravity direction), y the initial direction of the axis of the emerging flux rope and x the third orthonormal direction. The boundary conditions are kept similar for the seven parametric simulations for which the system is evolved with the same Lagrangian-remap code Lare3D (Arber et al. 2001). The numerical method is also presented in detail in L13 and L14. The simulations are performed in non-dimensional units, scaling that we are keeping in the present study. An example of a physical scaling is presented in L13 and L14.

The seven simulations differ by the properties of the initial ($t = 0$) background field that represents the pre-existing coronal field. In these simulations an arcade magnetic field is chosen which is invariant along the y -direction, i.e., parallel to the axis of the flux tube at $t = 0$. The magnetic field is given by

$$\mathbf{B}_{Arc} = \nabla \times \mathbf{A}_{Arc}, \text{ where} \quad (1)$$

$$\mathbf{A}_{Arc} = B_d \left(0, \frac{z - z_d}{(x^2 + (z - z_d)^2)^{3/2}}, 0 \right) \quad (2)$$

with z_d the depth at which the source of the arcade is located and B_d a signed quantity controlling the orientation (the sign of B_d) and the magnetic strength ($|B_d|$) of the arcade. This magnetic field has an asymptotic decay of $1/z^3$. The arcade is initially perpendicular to the convection zone flux tube's axis, and thus is aligned along the same axis as the twist component of the flux tube's magnetic field. The seven parametric simulations presented in the paper correspond to seven different values of B_d given on Table 1. The simulations with a positive B_d , which led to a stable configuration, have been presented in L13 while the simulations with a negative B_d , which induced an eruption, were presented in L14. The magnetic arcade strength $|B_d|$ has four different normalised values, i.e., [0,5,7.5,10], which respectively correspond to configurations with no arcade (ND), a weak arcade (WD), a medium arcade (MD) and a strong arcade (SD), following the notation of L13 and L14.

The simulation with $B_d = 0$ corresponds to the absence of coronal magnetic field. This simulation was presented both in L13 and L14 and used as a idealised reference simulation. In the real solar corona some small magnetic field is always present. L13 and L14 showed that the emergence of the flux rope in the magnetic-field-empty corona led to the formation of a stable magnetic structure in the solar corona. No eruptive behavior was observed. In the present manuscript this simulation is labeled "No Erupt ND".

The three simulations with $B_d > 0$ were presented in L13. For this configuration the direction of the arcade field is parallel to the direction of the upper part of the poloidal field of the twisted emerging flux rope. Fig. 1, top panels, presents the typical evolution of the system in the $B_d = 7.5$ case. In this set-up the flux rope emerges into a field in a configuration which is unfavourable for magnetic reconnection. As the initial flux rope

emerges, these simulations result in the formation of a new stable coronal flux rope. The flux rope, as with the no-arcade configuration, does not present any eruptive behavior and remains in the coronal domain until the end of the simulations. Hereafter, these three simulations are labeled "No Erupt WD", "No Erupt MD", "No Erupt SD" according to the strength of the magnetic arcade (respectively with $B_d = [5, 7.5, 10]$).

The three simulations with $B_d < 0$ were presented in L14. Fig. 1, bottom panel, displays the typical evolution of the system for the $B_d = -7.5$ simulation. For this configuration, the direction of the arcade field is antiparallel to the direction of the top of the twisted emerging flux rope. In this set-up the flux rope emerges into a coronal field for which the orientation is favourable for magnetic reconnection. Similarly to the non-eruptive simulations, the emergence results in the formation of a new coronal flux rope. However, in contrast to the $B_d \geq 0$ cases, here the new coronal flux rope is unstable and erupts. The secondary flux rope is immediately rising upward exponentially in time and eventually ejected up from the simulation domain. Once the ejection occurs, the remaining coronal magnetic field remains stable with no further eruption, even though residual flux emergence was still underway. In the following, these three simulations are labeled "Erupt WD", "Erupt MD", "Erupt SD" according to the strength of the magnetic arcade (respectively with $B_d = -[5, 7.5, 10]$).

Since observational studies are not yet able to provide information about the sub-photospheric magnetic field, we will only focus on analysing the atmospheric part of the flux emergence simulations of L13 and L14. The datasets analysed here thus correspond to the following subdomain: $x \in [-100, 100]$, $y \in [-100, 100]$, $z \in [0.36, 150]$. To comply with the constraints of our magnetic helicity estimation routines, the original irregular 3D grid of L13 and L14 is remapped on a 3D regular and uniform mesh, using trilinear interpolation. The grid used in the present study has a constant pixel size of 0.859 along all three dimensions. This corresponds to resp. [1.3,2] times the smallest horizontal and vertical grid spacing used in the original simulations in L13 and L14, located at the center of the simulation, and [0.33, 0.43] times its largest grid spacing, respectively. As will be discussed in Sect. 4.1 this interpolation deteriorates the level of solenoidality ($\nabla \cdot \mathbf{B} = 0$) of the magnetic field compared to the original data of L13 and L14.

The simulations last from $t = 0$ to $t = 200$. For all the simulations, before $t \sim 30$ the emerging flux rope rises in the convection zone, and the coronal field (i.e., the domain studied here), remains close to its initial stage. The rising flux ropes eventually reach the photospheric level, and in our datasets the emergence effectively starts around $t \sim 30$ (see Fig. 2). Eventually, in the eruptive simulations only, reconnection is observed, from $t \sim 120$. In the following the period $t \in [30, 120]$ will be referred as the pre-eruptive phase both for the eruptive and the non-eruptive simulations. An efficient proxy of eruptivity should be able to discriminate/distinguish the eruptive simulations from the non-eruptive ones *during* the pre-eruptive phase only.

Magnetic reconnection above the flux rope enables its ejection in the eruptive simulation. The flux rope eventually crosses the top boundary of our dataset (at $z = 150$) around $t \sim 150$. The period between $t \sim 120$ and $t \sim 150$ is named hereafter as the eruptive phase. The period with $t > 150$, until the final time analysed in this study at $t = 200$, is the post-eruptive phase. It should also be noted that the non-eruptive simulations have been carried out until $t > 450$. No eruptive behavior was sighted in that later phase of their evolution. At the end of the time range studied here, the non-eruptive simulations are thus still not in an erup-

tive stage. Hence during the post-eruptive phase, since both the non-eruptive and the eruptive simulations are in a stable state, an efficient eruptivity proxy should not be able to discriminate between them, in addition to being able to discriminate between them in the pre-eruptive phase.

3. Magnetic flux and energy evolution

Before analysing the differences in the magnetic helicity content in the simulations, it is important to first compare the evolution of the magnetic flux and magnetic energies, quantities which are typically used to characterise eruptive systems such as active regions. More frequent and powerful CMEs are known to originate from active regions with higher magnetic flux. As discussed in Sect. 1, they are theoretically expected and observationally found to have a relatively high non-potential energy, and hence have a greater reserve of energy to fuel the eruption.

3.1. Magnetic flux

The magnetic flux Φ , half the total unsigned flux, at the bottom boundary of the system is given by:

$$\Phi = \frac{1}{2} \int_{z=0} |\mathbf{B} \cdot d\mathbf{S}|. \quad (3)$$

Its evolution in time for the seven different simulations is represented in the top panel of Fig. 2. Because of the different values of the strength of the arcade, the initial magnetic flux $\Phi_{ini} \equiv \Phi(t = 0)$ in the seven simulations has different intensities. While the simulation with no arcade possesses no initial magnetic flux, the amount of Φ_{ini} in the other simulations is simply related to the arcade strength. As theoretically expected, one has $\Phi_{ini,MD}/\Phi_{ini,WD} = 3/2$ and $\Phi_{ini,SD}/\Phi_{ini,MD} = 4/3$. In the present simulation framework, it is obvious that the magnetic flux does not constitute a discriminative factor for eruptivity. The values of Φ for the eruptive and non-eruptive simulations are completely intertwined.

Because the simulation set-up consists of a flux rope emerging into a coronal field, it is interesting to plot the injected magnetic flux, here defined as the flux added to the pre-existing field. In the bottom panel of Fig. 2, we represent the injected magnetic flux, Φ_{inj} , defined, for each simulation, in reference to its initial magnetic flux, Φ_{ini} :

$$\Phi_{inj} \equiv \Phi - \Phi_{ini}. \quad (4)$$

The curves of Φ_{inj} show very strong similarities in terms of injected flux. This is expected since the very same magnetic structure is emerging in all seven runs. For all seven cases, the emergence starts around $t \sim 30$ and presents a very sharp increase until $t \sim 60$. During that period more than 80% of the magnetic flux is injected in the systems. In this first phase of the emergence, the curves are barely distinguishable from each other. The curves only start to be slightly different after $t \sim 65$. After that the magnetic injection increases moderately before reaching a plateau and slightly decreasing. The quantity Φ_{inj} is therefore not able to discriminate the eruptive behavior present in the different simulations. No distinctive signature of eruptivity is present in the curves of the magnetic fluxes in the pre-eruptive phase.

3.2. Magnetic energies

The magnetic energy being the central source of energy in active solar events, motivates us to present in Fig. 3, top left panel,

Table 1. Parametric simulations

Label	No Erupt SD	No Erupt MD	No Erupt WD	No Erupt ND	Erupt WD	Erupt MD	Erupt SD
B_d	10	7.5	5	0	-5	-7.5	-10
Arcade Strength	Strong	Medium	Weak	Null	Weak	Medium	Strong
Eruption	No	No	No	No	Yes	Yes	Yes

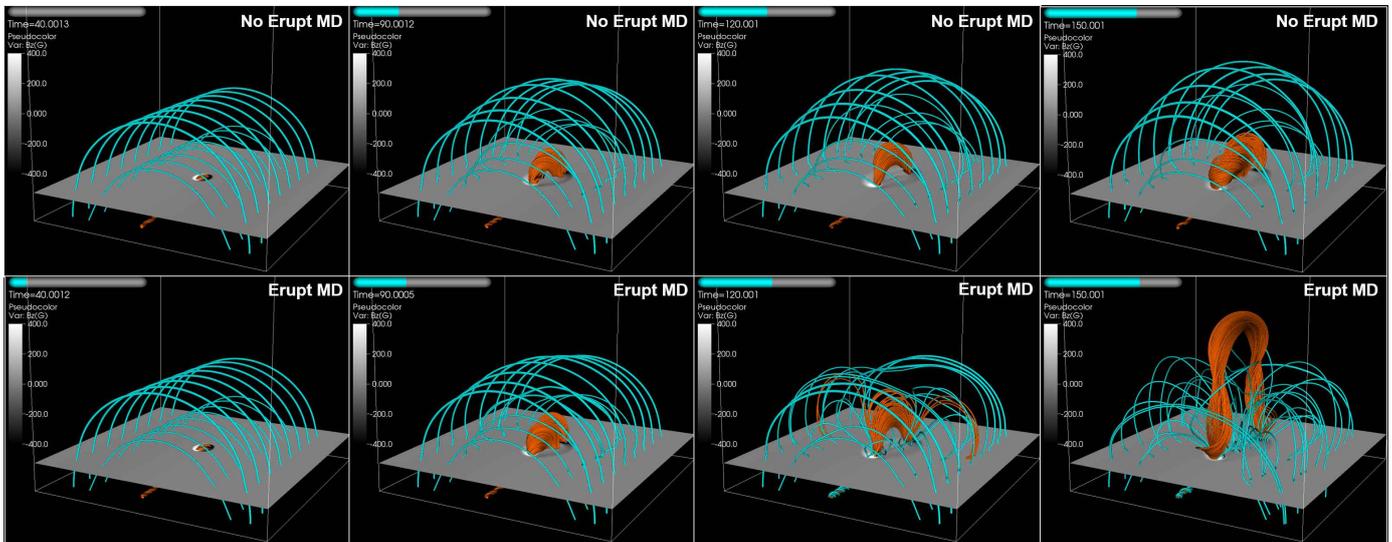


Fig. 1. Snapshots comparing the evolution of the systems in the eruptive (bottom row) and non-eruptive (top row) cases with medium arcade strength ($B_d = \pm 7.5$). The resp. [cyan, orange] field lines initially belongs to the [arcade, emerging flux rope]. The 2D horizontal cut displays the distribution at $z = 0$ of the vertical component of the magnetic field, B_z with a greyscale code. Only the volume above that boundary is considered in the present study.

the evolution of the magnetic energy E_{mag} for the different simulations. Similarly to Φ , because of the difference in the initial magnetic coronal field, E_{mag} does not constitute a pertinent criterion of eruptivity. For each simulation, we also plot the injected magnetic energy E_{inj} , defined relative to the initial value $E_{mag,ini} \equiv E_{mag}(t = 0)$:

$$E_{inj} \equiv E_{mag} - E_{mag,ini} . \quad (5)$$

As for Φ_{inj} , the evolution of E_{inj} for the different simulations in the initial phase of the emergence, between $t = 30$ and $t = 65$, presents extremely similar properties. One simulation is barely distinguishable from any other one. It is only once the system is erupting, for $t > 120$, that E_{inj} starts to present significant differences between the eruptive and non-eruptive cases. This is likely due to the ejection of the erupting current-carrying structure outside of the simulation domain. In any-case, this indicates that E_{inj} , similarly to Φ_{inj} , does not represent an efficient eruptivity criterion that would allow a forecast of the eruptions.

As discussed in Valori et al. (2013), the magnetic energy of a magnetic field with finite non-solenoidality ($\nabla \cdot \mathbf{B} \neq 0$), can be decomposed as:

$$E_{mag} = E_{pot} + E_{free} + E_{ns} , \quad (6)$$

where E_{pot} and E_{free} are the energies associated with the potential and current-carrying solenoidal contributions, respectively, and E_{ns} is the sum of the non-solenoidal contributions (see Eqs. (7,8) in Valori et al. 2013, for the corresponding expressions). The potential field is computed such as to match the normal component of the field on all six boundaries. In the case of a purely solenoidal field, $E_{ns} = 0$ in accord with the Thompson theorem. However, since numerical datasets never induce a perfectly null divergence of \mathbf{B} , a finite value of E_{ns} is generally

present. Unlike other physically meaningful energies, E_{ns} is a pseudo-energy quantity which can be positive or negative.

The different values of the decomposition of energy are plotted in Fig. 3. While the potential energy presented in the middle panel of Fig. 3 does not constitute an interesting criterion for eruptivity, its evolution is interesting for understanding the energy accumulation. Initially the system is fully potential, i.e., $E_{free} = 0$ and $E_{mag}(t = 0) = E_{pot}(t = 0)$. As the flux starts to emerge, the potential energy increases due to the modification of \mathbf{B} at the 6 side boundaries of the domain. While at the very beginning of the emergence, for $t \in [30, 40]$, the potential field of the eruptive and non-eruptive simulations shows a similar increase, the non-eruptive simulations possess a significantly higher potential energy compared to the eruptive simulation done at equivalent arcade strengths. At $t = 80$, the non-eruptive simulation of a given $|B_d|$ owns about 1.2 times more potential energy than its counterpart eruptive run. This further confirms that the potential energy, and its relative accumulation, cannot constitute a good eruptivity criterion.

We also notice that a large part of the injected magnetic energy, E_{inj} , is comprised of the increase in the potential energy, although not the majority. Before $t \sim 100$, the potential energy represents about 1/3 of the accumulated total magnetic energy. This shows, as in Pariat et al. (2015), that taking E_{inj} as a proxy for the free magnetic energy E_{free} , as is frequently done, can lead to substantial errors, and that properly computing the energy composition of Eq. (6), using the full boundary information, is an important step of any proper energy budget analysis.

The free magnetic energy, E_{free} , is a fundamental quantity in solar eruption theory. As the primary energy tank for all the dynamics of the phenomena developing during a solar eruption, its estimation is a main focus of solar flares studies (Tziotziou

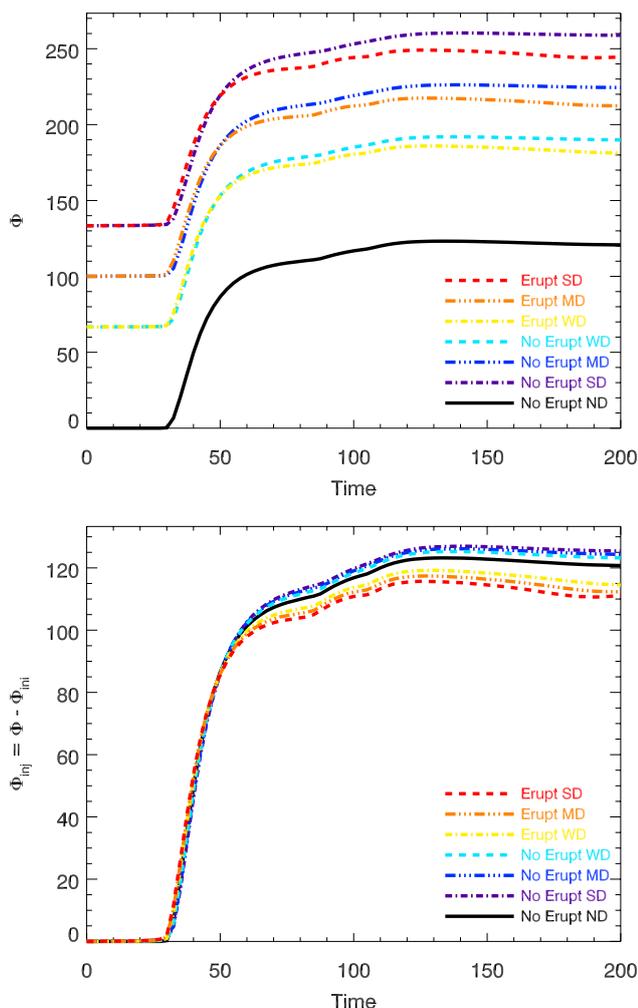


Fig. 2. Evolution of the absolute magnetic flux (Φ , top panel) and of the injected magnetic flux ($\Phi_{inj} \equiv \Phi - \Phi(t=0)$, bottom panel) in the system for the 7 parametric simulations. The non-eruptive simulation without surrounding field (No Erupt ND) is plotted with a continuous black line. The non-eruptive simulations with resp. [strong, medium, weak] arcade strength, labelled [No Erupt SD, No Erupt MD, No Erupt WD], are plotted respectively with a [purple dot dashed, blue dot-dot-dot-dashed, cyan dashed] line. The eruptive simulations with resp. [strong, medium, weak] arcade strength, labelled resp. [Erupt SD, Erupt MD, Erupt WD], are plotted respectively with a [red dashed, orange dot-dot-dot-dashed, yellow dot dashed] line.

et al. 2012, 2013; Aschwanden et al. 2014). The bottom right panel of Fig. 3 presents the time variations of E_{free} for the seven parametric simulations studied here. As for Φ_{inj} and E_{inj} , E_{free} presents a relatively similar dynamic for the seven simulation in the first half of the simulation. Unlike previous criteria, the curves of E_{free} for the eruptive simulations tend to be slightly higher than for the non-eruptive ones. The additional free magnetic energy remains, however, weakly higher. Interestingly, after $t \sim 100$ the values of E_{free} for the eruptive simulations tend to decrease and become lower than the ones of the non-eruptive simulations. This behavior confirms the common theoretical understanding that E_{free} is indeed a quantity tightly linked with the eruptive dynamics. Nonetheless, the highest values of E_{free} achieved are reached by the non-eruptive simulations in the post eruptive phase. Thus, while E_{free} certainly represents a neces-

sary condition for flares, it does not seem to be a significant sufficient condition for eruptivity.

The ratio of the free magnetic energy to the injected magnetic energy, E_{free}/E_{inj} , presents an interesting potential as an eruptivity proxy. Indeed, as shown in Fig. 4, the eruptive simulations are well distinguishable from the non-eruptive ones already in the pre-eruptive phase using this criterion. We note that the eruptive simulations present a higher ratio of E_{free}/E_{inj} than the non-eruptive simulations in the pre-eruptive and the eruptive phase, for $t \in [30, 130]$. After the eruption, this quantity decreases, which is expected from a good eruptivity proxy. However, we see that the non-eruptive simulations eventually present values of E_{free}/E_{inj} as high as the eruptive simulation in the initial phase. These simulations nonetheless do not present signs of eruptive behavior, not even beyond the time interval considered here. In addition, the values of E_{free}/E_{inj} are, for the eruptive simulations, only slightly superior to the non-eruptive ones. In practical cases, this criterion thus may not be very efficient.

Overall, even though the free magnetic energy, and its related quantities such as E_{free}/E_{inj} , are discriminative between the eruptive and non-eruptive simulations, it is only marginally so, in particular with regards to the criterion based on magnetic helicity that are discussed in Sect. 4.

Following Valori et al. (2013), we also compute the ratio $|E_{ns}|/E_{mag}$ which has been suggested as a meaningful estimation of the relative level of non-solenoidality present in the dataset. As discussed in Valori et al. (2013, 2016), this quantity is fundamental to establish the reliability of the magnetic helicity computation that we are presenting in Sect. 4. The bottom left panel of Fig. 3 shows that this solenoidality criterion remains relatively small throughout the different simulations, indicating a relatively good solenoidality of the data. While for most of the simulations, $|E_{ns}|/E_{mag} \approx 6\%$ before $t \sim 30$, the values quickly go below 2% for $t > 45$. It should be noted that the levels of non-solenoidality present in the data studied in this paper are very likely much higher than the one in the original datasets studied by L13 and L14. Indeed the interpolation performed to remap the data on a uniform grid from the original staggered grid have likely significantly increased the divergence of \mathbf{B} . Even then, as shown in Valori et al. (2016), the low fraction of $|E_{ns}|$ present here nonetheless ensures a good level of confidence of the magnetic helicity measurements.

4. Magnetic helicity evolution

4.1. Relative magnetic helicity measurements

The classical magnetic helicity (Elsasser 1956) of a magnetic field \mathbf{B} studied over a fixed fully-closed volume \mathcal{V} is gauge invariant only when considering a volume bounded by a flux surface, i.e., a volume whose surface $\partial\mathcal{V}$ is tangential to \mathbf{B} . In most practical cases, as in the present study, the studied volume surface is threaded by magnetic field. Following the seminal work of Berger & Field (1984), we therefore track here the evolution of the relative magnetic helicity. For relative magnetic helicity to be gauge invariant, the reference field must have the same distribution of the normal component of the studied field \mathbf{B} along the surface. A classical choice, adopted here, is to use the potential field \mathbf{B}_p as the reference field (see Prior & Yeates 2014, for a possible different class of reference field). As in Valori et al. (2012), we use here the definition of relative magnetic helicity from Finn & Antonsen (1985):

$$H_V = \int_{\mathcal{V}} (\mathbf{A} + \mathbf{A}_p) \cdot (\mathbf{B} - \mathbf{B}_p) d^3V. \quad (7)$$

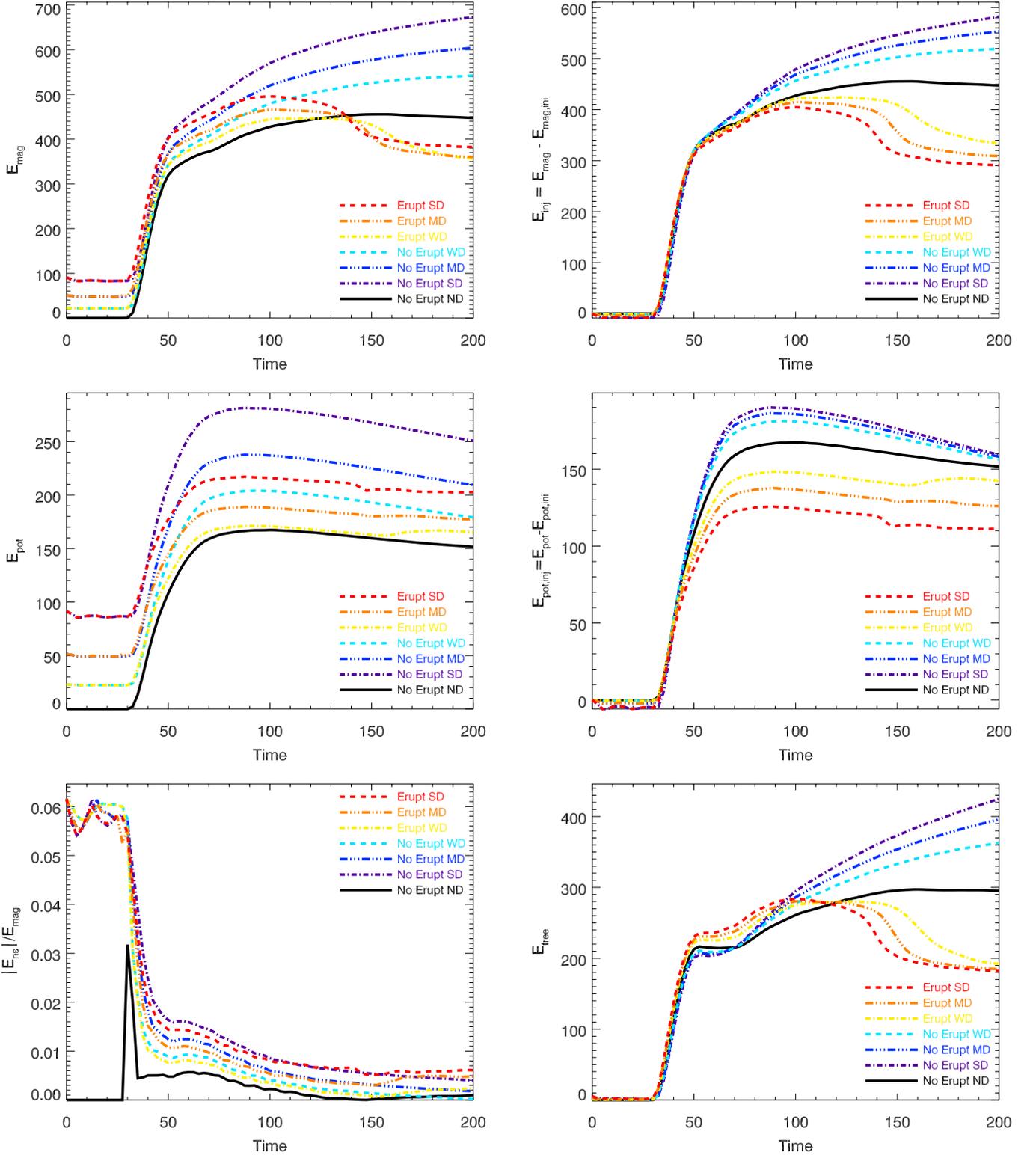


Fig. 3. Energy evolution in the system for the 7 parametric simulations: total magnetic energy (E_{mag} , top left panel), injected magnetic energy ($E_{inj} \equiv E_{mag} - E_{mag}(t=0)$), top right), Potential magnetic energy (E_{pot} , middle left), potential energy variation ($E_{pot} - E_{pot}(t=0)$), middle right), ratio of the artefact non-solenoidal energy to the total energy ($|E_{ns}|/E_{mag}$, bottom left) and free magnetic energy (E_{free} , bottom right). The labels are similar to Fig. 2.

with \mathbf{A} and \mathbf{A}_p the vector potential of the studied and of the potential fields: $\mathbf{B} = \nabla \times \mathbf{A}$ and $\mathbf{B}_p = \nabla \times \mathbf{A}_p$, respectively. Given the distribution of the normal component on the full surface $\mathbf{B} \cdot d\mathbf{S} = \mathbf{B}_p \cdot d\mathbf{S}$, the potential field is unique. Independently

of the time evolution of the magnetic system, this quantity is gauge invariant by definition. Even though the reference potential field may vary in time, along with possible evolution of the flux distribution on $\partial\mathcal{V}$, as demonstrated in Valori et al. (2012), it

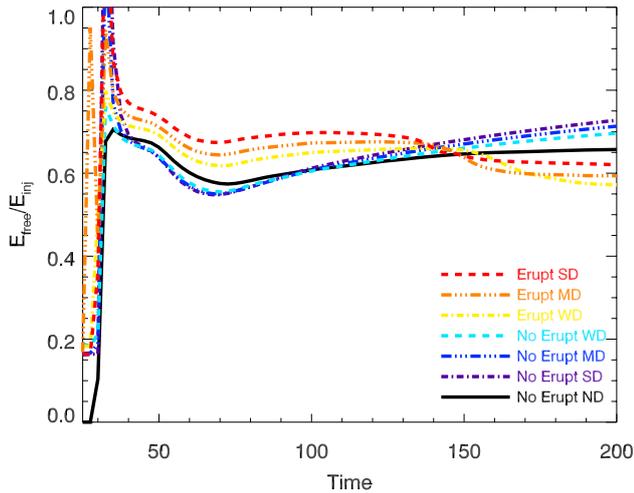


Fig. 4. Time evolution of the ratio of the free magnetic energy, E_{free} to the injected magnetic energy $E_{inj} \equiv E_{mag} - E_{mag}(t=0)$, for the seven parametric simulations. The labels are similar to Fig. 2.

is possible and physically meaningful to compute relative magnetic helicity (called magnetic helicity hereafter) and track it in time in order to characterise the evolution of a magnetic system.

Since magnetic helicity is an extensive quantity which scales with the square of a magnetic flux, it is of interest to study an intensive helicity-based quantity. In the following we will use the normalised helicity, \tilde{H} , given by the ratio of H_V to the square of the injected bottom-boundary magnetic flux, Φ_{inj} at the same time:

$$\tilde{H} = H_V / \Phi_{inj}^2. \quad (8)$$

In the case of a uniformly twisted flux tube, the normalised helicity would correspond to the number of turns of the magnetic field. The observational properties of this normalised helicity in the solar context have been reviewed in Démoulin & Pariat (2009).

A possible decomposition of relative magnetic helicity from Eq. (7) has been given by Berger (2003):

$$H_V = H_j + 2H_{pj} \quad \text{with} \quad (9)$$

$$H_j = \int_{\mathcal{V}} (\mathbf{A} - \mathbf{A}_p) \cdot (\mathbf{B} - \mathbf{B}_p) d\mathcal{V} \quad (10)$$

$$H_{pj} = \int_{\mathcal{V}} \mathbf{A}_p \cdot (\mathbf{B} - \mathbf{B}_p) d\mathcal{V} \quad (11)$$

where H_j is the classical magnetic helicity of the non-potential, or current carrying, component of the magnetic field, $\mathbf{B}_j = \mathbf{B} - \mathbf{B}_p$ and H_{pj} is the mutual helicity between \mathbf{B}_p and \mathbf{B}_j . The field \mathbf{B}_j is contained within the volume \mathcal{V} so it is also called the closed field part of \mathbf{B} . Because \mathbf{B} and \mathbf{B}_p have the same distribution on $\partial\mathcal{V}$, not only H , but also both H_j and H_{pj} , are theoretically independently gauge invariant. An alternative and widespread decomposition splits helicity into self, potential and mixed terms, see e.g., Eqs. (11-13) in Pariat et al. (2015). However, since the terms in that decomposition are not individually gauge-invariant, their separate evolution is devoid of any physical meaning, and hence it is not suitable for our purpose, and it is not considered here.

In order to estimate H_V , H_j and H_{pj} the quantities \mathbf{B}_p , \mathbf{A}_p , and \mathbf{A} must be derived. The effective computation of the potential

vectors requires the choice of a system of gauges. Several methods to compute H in a cartesian cuboid system have been developed in the recent years, using different choices of gauges and/or numerical approaches. Valori et al. (2016) presents a review of these methods and a benchmark of their efficiency. They found that as long as the studied field was sufficiently solenoidal, the derived helicity was overall consistent. Two of the simulations used in the present study have actually been used as test-cases of the benchmarking.

In the present study, we adopt the method of Valori et al. (2012) for the computation of the relative magnetic helicity. The potential vectors are computed using the DeVore gauge (DeVore 2000), i.e., $A_z = 0$. This method actually allows us to compute helicity with different sets of gauges (cf. Appendix A) as well as allowing us to determine the quality of the helicity conservation in the numerical domain (cf. Appendix B).

While H_V , H_j and H_{pj} are theoretically gauge invariant for purely solenoidal fields, the finite level of solenoidality of \mathbf{B} , inherently present in any discretised dataset, induces a certain gauge dependence of the helicities (Valori et al. 2016). In Sect. 3.2, we noted that the value of $|E_{ns}/E|$ lies at $< 6\%$ and $< 1\%$ depending on the phase of the emergence. The computation of relative helicity with different gauge sets allows us to control the impact of the finite non-solenoidality of the data on the helicity estimation. These tests are presented in Appendix A. Thanks to the relatively low level of $|E_{ns}/E|$, we find that the gauge invariance of the helicity quantities is well verified, with measurement errors on the helicity quantities on the order of 5%. Such a tolerance can be taken as an error on the helicity curves presented here. The conclusions that are drawn from our study are not affected by such an error.

4.2. Total magnetic helicity evolution comparison

The comparative evolution of magnetic helicity for the seven different flux emergence simulations is presented in Fig. 5. In all cases, the helicity presents a smooth increase in time as helicity is injected in the system thanks to the continuous emergence. In the eruptive cases, the ejection of the flux rope from the simulation domain is associated (very weakly in the strong arcade case) with a small decrease of H_V after $t \sim 150$. It is worth noting that, compared to the injection of magnetic flux and energy, the helicity accumulates much more smoothly and slowly. While more than 70% of Φ_{inj} and 50% of E_{free} is injected in the system between $t = 30$ and $t = 50$, only 10% of H_V has been injected in the system during that period. In all the flux emergence simulations studied here, the helicity injection is thus partly delayed compared to these other quantities. This delay between the magnetic flux increase and the helicity accumulation agrees with the trend noted in observational studies of active region emergence (Jeong & Chae 2007; Tian & Alexander 2008; Liu & Schuck 2012).

The first significant result is that, unlike for magnetic flux and magnetic energy, magnetic helicity presents very significant differences between the different simulations immediately after the very start of the emergence. Each simulation is easily distinguishable from the others as early as $t \sim 30$. The magnetic helicity is much more of a discriminant than the energies and accumulated magnetic flux are. Magnetic helicity is thus able to characterise very well the magnetic configuration, as it depends not only on the strength of the surrounding field, but also on its orientation relative to the emerging flux rope.

As discussed in the Introduction, a large absolute value of the total magnetic helicity has been frequently suggested as a potential proxy for flare eruptivity. In the framework of the present

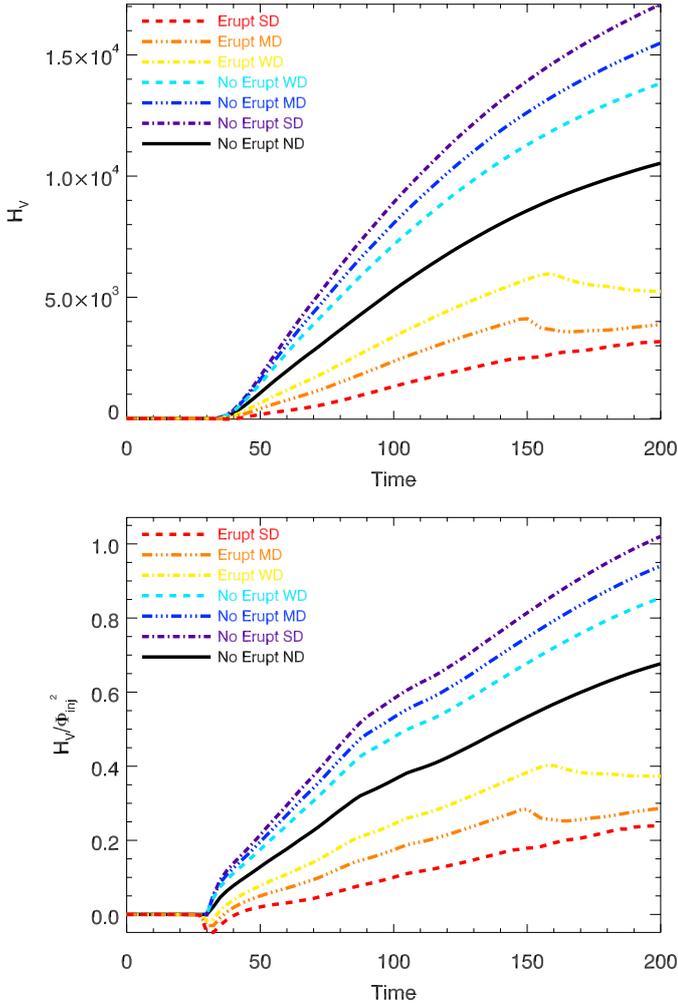


Fig. 5. H_V (top panel) and \tilde{H} (bottom panel) evolution for the seven parametric simulations. The labels are similar to Fig. 2.

simulation, we however notice that this is not the case. The top panel of Fig. 5 shows that the non-eruptive simulations all have a total absolute helicity $|H_V|$ several times higher than the eruptive one. Similarly the normalised helicity (Fig. 5, bottom panel) presents higher values for the non-eruptive cases. Our results indicates that a large value of $|H_V|$ cannot be used as a criterion for eruptivity.

Looking at the influence of the arcade field strength, we note that there is an opposite behavior for the eruptive and non-eruptive simulation when it comes to the total helicity. For the non-eruptive simulations, the stronger the arcade field, the greater the total helicity $|H_V|$ (and \tilde{H}), while for the eruptive cases the strength of the arcade and the intensity of $|H_V|$ are anti-correlated.

The origin of these behaviors can be first explained by the fact that the weaker the arcade strength, the closer the system is to the no-arcade case (an infinitely weak arcade would effectively correspond to an absence of arcade field). This explains why, given the orientation of the arcade, the curves of H_V and \tilde{H} tends to converge to the no-arcade case as the arcade strength decreases: the curves of \tilde{H} and H_V thus tends to become lower or higher, respectively for the non-eruptive and eruptive cases, as the arcade strength becomes lower.

This however does not explain why the orientation of the arcade leads to a higher H_V in the non-eruptive case and lower one for the eruptive simulations. This dependence originates from the fact that, unlike most quantities, magnetic helicity is intrinsically non-local (Berger & Murdin 2000). When the flux rope is emerging, it not only advects its own helicity, but also instantaneously exchanges helicity with the surrounding magnetic field. As we will show in the following simplified toy model, the difference of H_V between the different simulations are directly marked by the mutual helicity shared by the emerging flux rope and the arcade field.

In the case of a system formed by two closed flux tubes, the total helicity is the sum of the proper helicity contained in each of the flux tubes, their self helicity, plus the helicity shared between the flux tubes, their mutual helicity (cf. e.g. Berger 1984, 2003). In an analog toy model, one can theoretically decompose the helicity of the present system between the self helicity of the emerging flux rope, $H_{S,FR}$, the self helicity of the arcade field, $H_{S,Arc}$, and the mutual helicity shared between the emerging flux rope and the arcade field, H_{mut} :

$$H_V = H_{S,Arc} + H_{S,FR} + H_{mut} . \quad (12)$$

Initially, the flux rope not having yet emerged, the helicity of the system is solely given by the helicity of the arcade field and is expected to be null since the system is initially quasi-potential in the coronal domain (cf. Fig. 3), independently of the strength of the arcade field: $H_V(t=0) = H_{S,Arc} \simeq 0$. This is confirmed by the measured values of helicity at the beginning of the simulations in all cases. Furthermore, the simulation with no-arcade field contains no mutual helicity. The values of H_V in that case should roughly represent the evolution of the self helicity of the flux rope field: i.e., $H_{S,FR}(t) \sim H_{V, \text{No Erupt ND}}(t)$.

The differences in the curves of H_V should therefore originate from the difference in H_{mut} . In the case of two closed curved flux ropes, their mutual helicity is equal to the product of their magnetic flux weighted by their Gauss linking number (Berger & Field 1984; Berger & Murdin 2000). Depending on the relative orientation of the curves, the linking number can either be positive or negative. In the present simulation, it is reasonable to argue that $|H_{mut}|$ will be proportional to the product of the flux of the magnetic arcade, Φ_{Arc} , and the flux of the magnetic flux rope Φ_{FR} . The flux of the emerging flux rope is roughly constant between the simulations, and almost exactly in the initial phase of the emergence, before $t \sim 65$ as shown by the evolution of Φ_{inj} in Fig. 2. For each simulation, the flux of the arcade is directly given by the values of the flux initially, i.e., $\Phi_{Arc} = \Phi_{ini}$. The sign of the mutual helicity depends on the relative orientation of the arcade and the emerging flux rope: when the arcade and the axial field of the flux rope have a positive crossing, for the non-eruptive simulation, they should have a positive mutual helicity, while they should present a negative mutual helicity when the magnetic field orientation between the two has a negative crossing, for the eruptive cases. The helicity in the system should thus follow the relation:

$$H_V - H_{V, \text{No Erupt ND}} \equiv H_D \sim H_{mut} \sim \pm \xi \Phi_{Arc} \quad (13)$$

with ξ a constant of proportionality, and where the plus (minus) sign applies to the non-eruptive (eruptive) cases, respectively.

Qualitatively, this toy model predicts that the eruptive simulation should have a lower H_V than the no arcade case while the non-eruptive one should have greater values. In addition, the stronger the arcade field is, the further away H_V is from the no-arcade case. Quantitatively one also finds

very good agreement between the equation Eq. (13) predicted by this simple toy model and the actual values of H_V measured. During the main part of the emergence, when the system is not too affected by the ejection of magnetic field, one measures at $t = 75$ that $H_{D,\text{No Erupt MD}}/H_{D,\text{No Erupt WD}} = 1.46$, and $H_{D,\text{Erupt MD}}/H_{D,\text{Erupt WD}} = 1.51$, while our toy model theoretically predicts that these ratios should be equal to the ratio of the arcade strength between the medium arcade and the weak arcade case, i.e., $\Phi_{\text{ini,MD}}/\Phi_{\text{ini,WD}} = 3/2$. Similarly, one has $H_{D,\text{No Erupt SD}}/H_{D,\text{No Erupt MD}} = 1.30$, and $H_{D,\text{Erupt SD}}/H_{D,\text{Erupt MD}} = 1.3$, which should be equal to $\Phi_{\text{ini,SD}}/\Phi_{\text{ini,MD}} = 4/3$ according to our toy model. In the case of the non-eruptive simulations, the agreement improves as the emergence further develops.

The excellent agreement between these values demonstrates the importance of the mutual helicity between the emerging structure and the surrounding field, which, being added to or subtracted from the helicity advected by the emerging flux rope, significantly modifies the total amount of helicity. Even though in the seven simulations the emerging structure is exactly the same, the helicity budget is profoundly modified by the surrounding field. This highlights the importance of the surrounding environment when considering the budget of magnetic helicity in flux emergence regions, unlike with more classical quantities such as energies.

While self and mutual helicity are useful theoretical concepts, they are in practice very difficult to use. The distinction between the emerging field and the arcade field is strongly subjective. When only considering a unique snapshot of one of the simulations, it is very difficult to objectively disentangle these two structures. It is even more difficult to directly compute each contribution. It is only thanks to the combined seven parametric simulations that we are able to compute the respective self and mutual contributions in the present study. Unlike H_j and H_{pj} , we do not believe that it is generally possible to estimate these quantities from general datasets.

4.3. Current-carrying magnetic helicity evolution comparison

We have seen that while the total magnetic helicity H_V is very discriminative of the different parametric simulations, its use to predict the eruptive behavior is limited, since for a given magnetic flux injection, the eruptive simulations possess a lower total amount of helicity. We will now discuss the evolution of the terms H_j and H_{pj} decomposing the magnetic helicity (cf. Eq. (9)) and show that they constitute a very promising criterion of eruptivity.

The time evolution of H_j and H_{pj} for the seven parametric simulations is presented in Fig. 6. We note that while the curves of H_{pj} present a relative distribution very similar to H_V between the different simulations, the curves of H_j differ significantly. Regarding H_{pj} , the non-eruptive simulations have an evolution similar to each other and similar to H_V both in shape and in amplitude. The differences in the evolution of H_{pj} for the eruptive simulations is more marked. While the weaker-arcade-strength case presents a relatively smooth increase, the stronger-arcade case displays slightly negative values during the first part of the emergence until $t \sim 100$ with then some increase. As for H_V , no specific behavior before the onset of the eruption is noticeable in the evolution of H_{pj} .

The second main result of our study is that the curves of H_j are on the contrary strongly marked by the eruptive behavior (Fig. 6, bottom panel). We note that the non-eruptive and eruptive simulations present two very distinct groups. The simula-

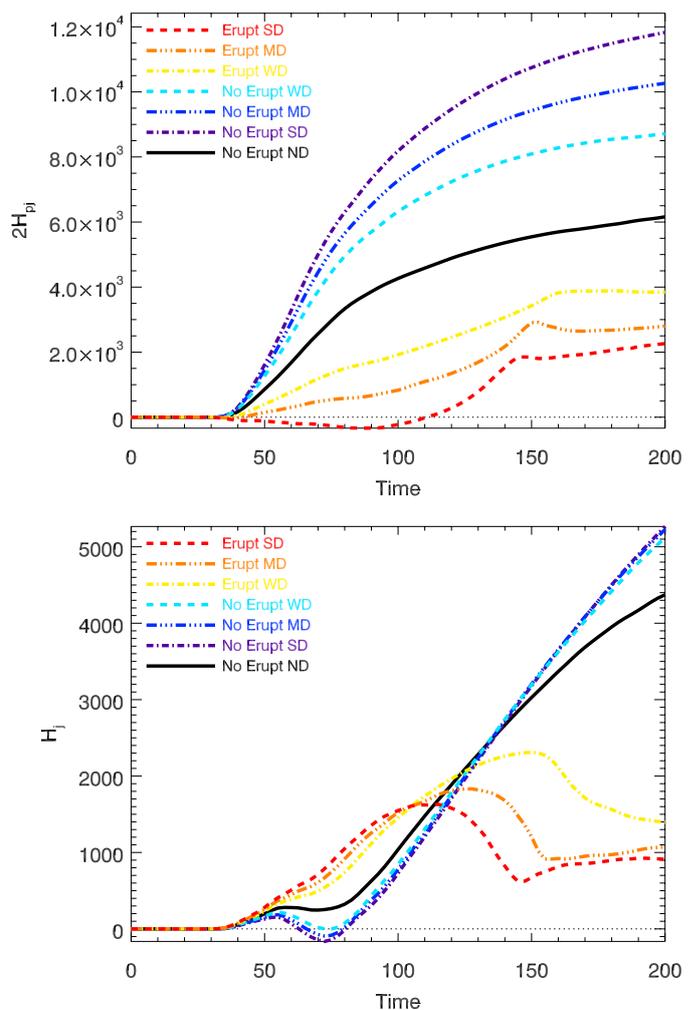


Fig. 6. Time evolution of $2H_{\text{pj}}$ (top panel), and H_j (bottom panel) for the seven parametric simulations. The labels are similar to Fig. 2

tion without surrounding magnetic field separates the two groups of simulations. For that simulation, after a slight increase for $t \in [30, 55]$, H_j presents a plateau until $t \sim 75$, before presenting a slow and steady increase. The three non-eruptive simulations with a surrounding magnetic field present a relatively similar behavior to each other. They are tightly grouped and are similar to the no arcade case. Instead of presenting a plateau between $t = 30$ and $t = 50$, H_j decreases, even reaching negative values before steadily increasing after $t \sim 75$. Unlike for H_V and H_{pj} , the field strength of the arcade does not seem to significantly influence the evolution of the values of H_j , although we note that the weak arcade curve is the one closest to the no arcade case.

The eruptive simulations are tightly grouped with each other. Unlike the non-eruptive simulations, they present a quasi-steady increase in the first half of the simulation, before $t \sim 100$. The curves then eventually reach a maximum, decrease and then stay relatively constant. While in the first part of the simulation, the curves only differ slightly in intensity, the timing of the maximum and the subsequent evolution is strongly influenced by the arcade field strength. The occurrence time of the maximum is anti-correlated with the strength of the arcade. The stronger the arcade the earlier the peak of H_j . This is likely correlated to the difference in the eruption time for the different eruptive simulations. As noted in Figure 12 of L14, the stronger the arcade, the

earlier the flux rope moves and is eventually ejected, leaving the domain. The differences between the curves of H_j for the eruptive simulations have the following explanation: the stronger the external arcade, the larger is the flux available for reconnection. For the emerging flux rope to erupt, the shell of stabilizing field surrounding it must be removed. Since the emergence timescale is dictated by the same photospheric evolution, more flux is available for reconnection for a given flux-rope emergence rate, the faster is the peeling of the outer shell, and the earlier is the time of eruption. Hence, the stronger is the external dipole field, the earlier is the start of the eruption. As a corollary, the longer it takes to erupt, the more of the flux rope emerges and, therefore, the higher is the maximum of H_j that can be reached.

While the eruptive simulations all display a higher value of H_j in the initial phase of the flux emergence, before $t \sim 100$, overall, the non-eruptive simulations are the ones which present the highest values of H_j in the later time of the evolution. Hence, the value of H_j alone, while being clearly affected by the eruptive behavior, cannot directly be used as an eruptivity criterion. The high value of H_j in the second phase of the simulations would otherwise suggest that the non-eruptive simulations could become unstable which does not agree with the dynamics observed at the end of these simulations. The situation is somehow similar to what was found for the free magnetic energy (cf. Fig. 3, bottom right panel), with the difference being that helicity spreads the curves farther apart, i.e., discriminates better between the different cases.

If H_j itself does not constitute an obvious eruptivity criterion, it nonetheless represents a significant portion of the total helicity of the system for the eruptive simulations, for example as can be seen in Fig. A.1 for the medium arcade case. Actually, the fraction of H_j to the total helicity is a key distinction between, first, the eruptive and, second, the non-eruptive simulation, but also between the pre-eruptive phase and the post-eruptive phase of the eruptive simulations.

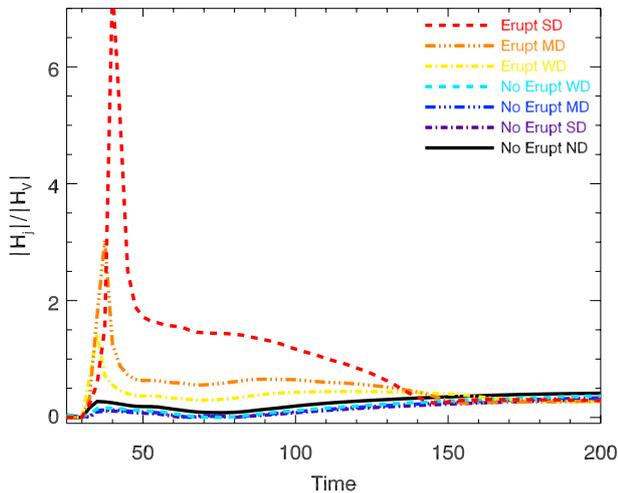


Fig. 7. Time evolution of the helicity ratio $|H_j|/|H_V|$ for the seven parametric simulations. The labels are similar to Fig. 2.

Fig. 7 presents the ratio of $|H_j|$ to the total helicity $|H_V|$ for the seven parametric simulations. Because these curves correspond to a ratio, and because H_V is roughly null before $t \sim 30$ (magnetic flux only increases from that time on), we only plot values after that time in order to remove the spurious values resulting from the division by an infinitely small value. The non-eruptive curves are approximately constant throughout the sim-

ulation, with values that do not exceed 0.4. For the four non-eruptive simulations, H_j always remains a minor contributor to H_V .

On the contrary, the eruptive simulations all present high values of $|H_j|/|H_V|$ during the first phase of the simulation. Immediately after the start of the emergence at $t \sim 30$, the curves present a very fast rise, with a peak between $t = 35$ and $t = 40$. The values of $|H_j|/|H_V|$ even exceed 1, indicating an opposite sign between H_j and H_{pj} . Helicity injection with opposite sign through the photospheric surface has been reported in several observed cases of eruptive active regions (Park et al. 2012; Vemareddy et al. 2012; Vemareddy & Démoulin 2016). Here, the level of the peak appears to be correlated with the strength of the arcade. It should be noted however that the values of H_V are still very small at these times, amounting to less than 2% of the helicity eventually injected. After this initial peak, the values of $|H_j|/|H_V|$ remain high, relative to the non eruptive simulations, with values globally above 0.45. The level of $|H_j|/|H_V|$ is also directly correlated with the strength of the arcade field, the stronger arcade field presenting a larger ratio than the medium arcade. The lower arcade presents the smaller ratio among the eruptive simulations, although still markedly higher than the non-eruptive simulation, with values 2 to 4 times higher than the no-arcade case, and more than 5 times higher than with the other stable runs. The curves remains relatively constant for some time, eventually decreasing after $t \sim 85$ and finally, after $t \sim 135$ joining the group of the curves of the non-eruptive simulations.

The values $|H_j|/|H_V|$ are not only higher for the eruptive simulations compared to the non-eruptive ones, but they are only so during the pre-eruptive and eruptive phase of the simulations. In the post eruption phase of the eruptive simulations, when the system is not eruptive, these values are back to a low value, below 0.4, typical of the non eruptive simulations. Furthermore, during the eruptive phase, the ratio $|H_j|/|H_V|$ is also markedly higher when the strength of the surrounding arcade is higher which, as noted by L14, is also related to a higher propensity for the system to erupt. Indeed, the higher arcade strength was associated with an earlier eruption of the system and a larger amount of reconnection. The ratio $|H_j|/|H_V|$ thus appears as a very interesting criterion for qualifying and possibly quantifying the eruptivity of a system in solar like conditions.

5. Eruptivity criteria comparison

In the previous sections, several scalar quantities that can characterise the magnetic data have been computed and their evolution analysed and compared for the seven parametric simulations. Their ability to constitute a pertinent criterion of the eruptivity of the system has been qualitatively discussed. Only positive-defined quantities are been considered in the present study and constructed so that eruptivity could be possibly associated with a high value of these proxies. Let us note that such positive criteria can always be built by the use of opposite or the inverse and modulus functions. In order to quantify the quality of a good eruptivity proxy we compute different parameters that evaluate their distribution between the different simulations: the mean value, μ , and the relative standard deviation, C_v , of a given quantity at a given time for all seven simulations; the mean values, μ_{Erupt} and $\mu_{\text{No Erupt}}$ only considering respectively the eruptive/non-eruptive simulations, and their ratio $\eta = \mu_{\text{Erupt}}/\mu_{\text{No Erupt}}$.

A quantity which is not able to distinguish between the different simulations will have C_v close to 0, as well as a value of η close to 1. This quantity thus does not possess the quality of a good proxy.

A high value of C_v indicates that this quantity is discriminating between the different simulations, although not necessarily their eruptive/non-eruptive character. Associated with a value of η close to 1, this means that this proxy is mostly sensitive to the strength of the surrounding arcade rather than the eruptivity.

A small value of η , close to 0, indicates that this quantity is significantly higher for the non-eruptive simulations. This generally means that this quantity is not a good proxy for eruptivity. On the contrary, a high value of η , associated with a large value of C_v , is what is required from a good eruptivity proxy since it indicates that this quantity tends to be significantly higher for the eruptive simulations compared to the non-eruptive one. For the eruptive simulations, a good eruptivity proxy should only have high values of C_v and η during the pre-eruptive and the eruptive phase and not during the post eruptive phase.

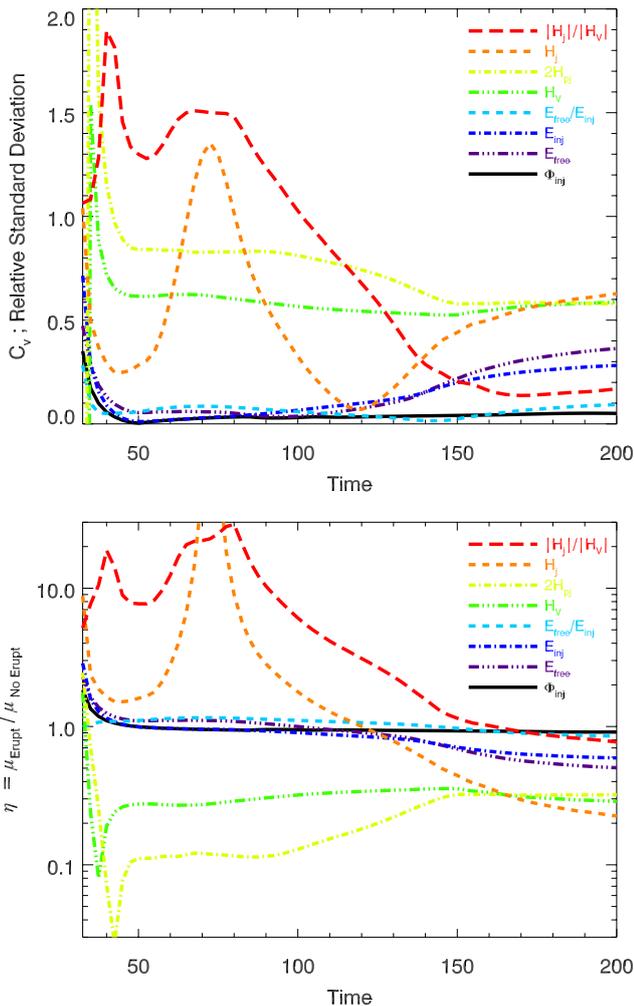


Fig. 8. Time evolution of the relative standard deviation, (C_v , top panel) and ratios of the mean eruptive and non-eruptive values, (η , bottom panel) of several potential criteria for eruptivity. The [black continuous, purple three-dot-dashed, blue dot-dashed, cyan dashed, green three-dot-dashed, yellow dot-dashed, orange dashed, red long-dashed] curve corresponds, respectively, to the quantity [Φ_{inj} , E_{inj} , E_{free} , E_{free}/E_{mag} , E_{free}/E_{inj} , H_V , $2H_{pj}$, H_j , $|H_j|/|H_V|$].

Fig. 8 presents the evolution of η and of the relative standard deviation between the different simulation. The curves of C_v show that the criteria based on magnetic flux and magnetic energy, namely, Φ_{inj} , E_{inj} , E_{free} , E_{free}/E_{mag} , E_{free}/E_{inj} are not able to distinguish between the different simulations. Their rel-

ative standard deviation remains very small, $< 20\%$, during the pre-eruptive phase, for $t \in [40, 70]$. On the contrary the magnetic helicity based proxies, H_V , H_j , H_{pj} , $|H_j|/|H_V|$ present high standard deviations indicating that the different simulations are strongly discriminating between these quantities. $|H_j|/|H_V|$ possesses the noticeable property of being very high in the pre-eruptive phase and in the eruptive phase, while decreasing in the post-eruptive phase.

The curves of η confirm again that the flux injection, Φ_{inj} , is extremely similar for the eruptive and the non-eruptive simulations. The mean value of the eruptive simulations is constantly almost equal to the mean value of the non-eruptive ones. This quantity thus does not present any quality looked for in an eruptive proxy. The same is true for E_{inj} . The total volume helicity H_V and H_{pj} , while displaying distinctive behavior between eruptive and non-eruptive, have very low values of η . This indicates that the eruptive simulation tends to have weaker values than the non-eruptive one. We observed in Sect. 4, that the highest values of these quantities were eventually reached by the non-eruptive simulations during the post-eruption phase. It is therefore not possible to define a threshold from these quantities which are thus poor eruptivity criteria.

The free magnetic energy, as well as E_{free}/E_{mag} and E_{free}/E_{inj} , shows a weak tendency to be higher for the eruptive simulation during their eruptive phase. E_{free} nonetheless has an η value that drops markedly below 1 in the post-eruptive phase. As shown in Fig. 3 the non-eruptive simulations present the highest values of E_{free} during that period. No threshold on E_{free} can thus be built in the present simulation framework. While being notably higher than 1 during the pre-eruptive phase, $\eta(E_{free}/E_{inj})$ decreases in the post-eruptive phase and becomes close to 1. Nonetheless, the value of η , of about 1.15 during the pre-eruptive phase, is not very high and thus may be of little practical use with real data. In addition, as already noted in Sect. 3.2, the non-eruptive simulations reach greater values of E_{free}/E_{inj} in the post-eruptive phase, even though no eruption is present (cf. L13). The ratio E_{free}/E_{inj} thus likely do not constitute a reliable proxy for eruption prediction.

The current magnetic helicity H_j presents a high value of η during the pre-eruptive phase, but then presents a low value during the post-eruptive phase, since the non-eruptive simulations present the highest values of H_j . Similarly to E_{free} , no threshold on H_j can be constructed that would enable the prediction of the eruptivity of our parametric simulation.

Finally, Fig. 8 confirms that the ratio $|H_j|/|H_V|$ is an extremely efficient proxy of eruptivity for the simulations. This quantity presents clear variations of C_v from the pre-eruption phase to the post-eruption phase. This quantity present a very high η , with values > 5 , in the pre-eruptive phase (for $t \in [30, 120]$). The values of $\eta(|H_j|/|H_V|)$ decrease during the eruptive phase and eventually become close to 1 during the non-eruptive phase, for $t > 150$, indicating that the eruptive and non-eruptive simulation are not distinguishable any more. This is expected since none of the simulations presents any eruptive behavior during that last phase. An eruptivity threshold can easily be built from the ratio of $|H_j|/|H_V|$. This quantity thus possesses a very strong potential to allow the prediction of solar eruptions.

6. Conclusion and discussion

In the present study, we have computed and compared the magnetic energy and helicity evolution of the coronal domain of seven parametric 3D MHD simulations of flux emergence, initially presented in Leake et al. (2013) and Leake et al. (2014).

These numerical experiments, while only modifying a unique parameter - the strength and direction of the background coronal field - led either to a stable configuration or to an eruptive behavior with the ejection of a CME-like magnetic structure. These simulations represent a particularly interesting dataset that enables us to search for eruptivity criteria.

Following the method of decomposition of the magnetic energy of Valori et al. (2013) and the method to compute relative magnetic helicity presented in Valori et al. (2012), we have computed different magnetic flux, energy and helicity based quantities throughout the evolution of the systems. As expected from the numerical set-up, we noted that all the simulations presented a quasi-similar injection of magnetic flux. We have found that unlike magnetic flux and energy, relative magnetic helicity very clearly discriminated between eruptive and stable simulations.

We have however found that the total magnetic helicity was not correlated with a stronger eruptive behavior. Non-eruptive simulations in fact presented a higher absolute value of the total magnetic helicity compared to the eruptive one. Using a toy-model we have shown that the non-eruptive simulation possessed self helicity of the emerging flux rope of the same sign as the mutual helicity between the emerging flux rope and the coronal background field. Eruptive simulations presented a lower total helicity because the self and mutual helicities were of opposite sign.

Our results thus confirm the ones from Phillips et al. (2005), stating that the total magnetic helicity is not a determining factor for CME initiation and that there might not be a universal threshold on eruptivity based on total magnetic helicity. We however argue against their conclusion that helicity in general is unimportant. Their set-up, similarly to our eruptive cases, present large helicities of opposite signs. The decomposition of helicity, if not its distribution, seems to be related to enhanced eruptive behavior. Using the helicity decomposition of the relative magnetic helicity in the current-carrying magnetic helicity, H_j and its counterpart $2H_{pj}$, introduced by Berger (2003), eruptive and non-eruptive cases present noticeably distinct behavior, with the eruptive simulations presenting significantly greater values of H_j during the pre-eruptive phase.

Comparing the different quantities in their capacity to efficiently describe the eruptivity status of the different simulations during their evolution, we noted that while the ratio of the free magnetic energy to the injected magnetic energy, E_{free}/E_{inj} is higher for the eruptive simulations, in their pre-eruptive phase only, it presents two drawbacks: the values of this quantity are only marginally higher (by a few %) for the eruptive simulation, compared to the non-eruptive ones, and the non-eruptive simulations can reach values of a similar amplitude to the eruptive case. The definition of an eruptivity threshold seems to be difficult to determine with such a quantity.

The quantity that appears as an excellent eruptivity proxy is the ratio of the current carrying magnetic helicity to the total helicity, $|H_j|/|H_V|$. This ratio is several times higher for the eruptive simulations compared to the non-eruptive simulations during the pre-eruptive phase, and similar for all simulations during the non-eruptive phase. None of the non-eruptive simulations present values of $|H_j|/|H_V|$ higher than 0.45 while all the eruptive simulations can reach values three times that threshold. The strong and medium arcade cases, which present the fastest eruptive behaviors, have values of $|H_j|/|H_V|$ higher than 0.45 for almost all their pre-eruptive phase. In the framework of the present parametric simulations, the quantity $|H_j|/|H_V|$ thus constitutes an excellent eruptivity proxy.

The use of parametric numerical simulations has enabled us to highlight the existence of a very promising proxy of eruptivity that could allow for improvement in our forecast of solar flares. One should however be conscious of the limits of our approach. The first limitation is of course the level of realism of the present simulations compared to the real Sun. In addition to the inherent limits of the MHD paradigm, two aspects of the numerical set-up may deviate from the conditions found in the Sun. The first is the relative intensity of the background field to the emerging flux rope: the magnetic flux of the background coronal arcade is of similar scale to the flux injected by the emerging flux rope. This is different from most large-scale active regions for which the initial field is notably smaller than the emerging flux. While non-dimensional, the scale of these simulations corresponds to a small scale emerging structure. In addition, the amount of helicity injected by the emerging structure is relatively high compare to typical solar values. Observations have shown that emerging active regions present normalised helicities \tilde{H} (cf Eq. (8)) comprised between 0.01 and 0.1 (LaBonte et al. 2007; Jeong & Chae 2007; Tian & Alexander 2008; Yang et al. 2009; Démoulin & Pariat 2009) which is several times smaller than the values obtained here. The emerging structure in the present simulation is thus injecting more twist in the coronal domain than typical active regions do. These limitations do not question the validity of our study and its results, but only its uncritical application to real data.

Furthermore, unlike most observational studies that, when testing eruptivity proxies, are incorporating active regions of various size, with magnetic fluxes ranging over orders of magnitude, our present dataset is limited to simulated active regions of quasi-identical size. Observational statistical studies show that larger active regions (in the sense of magnetic flux) have a greater eruptivity and flare productivity (e.g. LaBonte et al. 2007; Park et al. 2010; Tziotziou et al. 2012). The present study is thus not addressing this property. The eruptivity proxy $|H_j|/|H_V|$ may only be a determinant eruptivity criterion between active regions of similar magnetic flux. When comparing regions of different size, its ability may be reduced. Most observational studies indeed show that it is a combination of proxies that gives the best eruptivity criterion (Bobra & Couvidat 2015; Bobra & Ilonidis 2016). It is however interesting to note that $|H_j|/|H_V|$ is an intensive quantity and thus its capacity to predict eruptions may be independent from the active region size.

The final caveats that may limit the usage of $|H_j|/|H_V|$ is intrinsically related to our limited understanding of the properties of magnetic helicity. Relative magnetic helicity as defined by Eq. (7), is not a simply additive quantity (Berger & Murrin 2000). While all the computations done in the present study have been performed using the very same coronal domain, it is not guaranteed that the values of the ratio $|H_j|/|H_V|$ do not depend on the volume and boundaries of the studied domain. It is important to keep this in mind since it implies that the value of 0.45 discussed previously is likely to be only valid with regards to the present simulations given the size of the coronal domain chosen here. Further investigation are needed in order to better comprehend the properties of magnetic helicities.

Despite, these limitations, our study shows that magnetic helicity based quantities may be crucial in predicting solar flares. Because of its difficult estimation in observational cases, this class of quantity has been largely neglected in systematic studies searching for eruptivity proxies. Based on the present work, we argue that an additional effort should be carried out to test relative magnetic helicity based quantities against eruptive behavior

of observed active regions. We have demonstrated that the ratio $|H_j|/|H_V|$ appears to be a very promising intensive quantity.

The theoretical reasons explaining why this ratio is so efficient at describing the eruptive states of our parametric simulations still need to be fully explored. It is however interesting to note that H_j corresponds to a description of the current part of the field while the $|H_V|$, which is dominated by $2H_{pj}$, includes the relative properties of the current carrying field in relation to its surrounding field. In the framework of the torus instability (Kliem & Török 2006; Démoulin & Aulanier 2010; Zuccarello et al. 2015), the criterion of instability is also related to the relative properties of the current carrying structure with regards to its surrounding confining field. We here speculate that the ratio $|H_j|/|H_V|$ may somehow be related with the instability criterion of the torus instability. Further theoretical research should be engaged to further understand the link between the relative magnetic helicity decomposition and the eruptivity of a system.

Acknowledgements. EP acknowledge the support of the French Agence Nationale pour la Recherche through the HELISOL project, contract n° ANR-15-CE31-0001, as well as the FLARECAST project, funded by the European Union's Horizon2020 research and innovation programme under grant agreement n° 640216. JEL and MGL acknowledge support from the Chief of Naval Research and from the NASA Living with a Star program. The flux emergence simulations analysed in this work were performed under a grant of computer time from the US Department of Defense High Performance Computing program. GV acknowledges the support of the Leverhulme Trust Research Project Grant 2014-051. FPZ. is a Fonds Wetenschappelijk Onderzoek (FWO) research fellow (Project No. 1272714N). KD acknowledges support from the Computational and Information Systems Laboratory and the High Altitude Observatory of the National Center for Atmospheric Research, which is sponsored by the National Science Foundation. GV and EP thanks ISSI, its members and the participants of the ISSI international team on *Magnetic Helicity estimations in models and observations of the solar magnetic field* where this work has been discussed and commented.

References

- Arber, T. D., Longbottom, A. W., Gerrard, C. L., & Milne, A. M. 2001, *Journal of Computational Physics*, 171, 151
- Aschwanden, M. J., Xu, Y., & Jing, J. 2014, *The Astrophysical Journal*, 797, 50
- Barnes, G., Leka, K. D., Schrijver, C. J., et al. 2016, *The Astrophysical Journal*, 829, 89
- Baumann, G., Haugbølle, T., & Nordlund, Å. 2013, *The Astrophysical Journal*, 771, 93
- Berger, M. A. 1984, *Geophysical and Astrophysical Fluid Dynamics* (ISSN 0309-1929), 30, 79
- Berger, M. A. 2003, *Advances in Nonlinear Dynamical Systems: The Fluid Mechanics of Astrophysics and Geophysics*, 20030424, 345
- Berger, M. A. & Field, G. B. 1984, *Journal of Fluid Mechanics* (ISSN 0022-1120), 147, 133
- Berger, M. A. & Murdin, P. 2000, *Encyclopedia of Astronomy and Astrophysics*, 2403
- Bobra, M. G. & Couvidat, S. 2015, *The Astrophysical Journal*, 798, 135
- Bobra, M. G. & Ilonidis, S. 2016, *The Astrophysical Journal*, 821, 127
- Carlsson, M., Hansteen, V. H., Gudiksen, B. V., Leenaarts, J., & de Pontieu, B. 2016, *Astronomy and Astrophysics*, 585, A4
- Démoulin, P. & Aulanier, G. 2010, *The Astrophysical Journal*, 718, 1388
- Démoulin, P. & Pariat, E. 2009, *Advances in Space Research*, 43, 1013
- DeVore, C. R. 2000, *The Astrophysical Journal*, 539, 944
- Elsasser, W. M. 1956, *Reviews of Modern Physics*, 28, 135
- Emslie, A. G., Dennis, B. R., Shih, A. Y., et al. 2012, *The Astrophysical Journal*, 759, 71
- Falconer, D., Barghouty, A. F., Khazanov, I., & Moore, R. 2011, *Space Weather*, 9, 04003
- Falconer, D. A., Moore, R., Barghouty, A. F., & Khazanov, I. 2014, *Space Weather*, 12, 306
- Finn, J. H. & Antonsen, T. M. J. 1985, *Comments on Plasma Physics and Controlled Fusion*, 9, 111
- Guennou, C., Pariat, E., Leake, J. E., & Vilmer, N. 2017, in prep.
- Inoue, S., Hayashi, K., Magara, T., Choe, G. S., & Park, Y. D. 2014, *The Astrophysical Journal*, 788, 182
- Jeong, H. & Chae, J. 2007, *The Astrophysical Journal*, 671, 1022
- Jiang, C., Feng, X., Wu, S. T., & Hu, Q. 2012, *The Astrophysical Journal*, 759, 85
- Jing, J., Chen, P.-F. F., Wiegmann, T., et al. 2009, *The Astrophysical Journal*, 696, 84
- Jing, J., Tan, C., Yuan, Y., et al. 2010, *The Astrophysical Journal*, 713, 440
- Kliem, B. & Török, T. 2006, *Physical Review Letters*, 96, 255002
- Kusano, K., Bamba, Y., Yamamoto, T. T., et al. 2012, *The Astrophysical Journal*, 760, 31
- LaBonte, B. J., Georgoulis, M. K., & Rust, D. M. 2007, *The Astrophysical Journal*, 671, 955
- Leake, J. E., Linton, M. G., & Antiochos, S. K. 2014, *The Astrophysical Journal*, 787, 46
- Leake, J. E., Linton, M. G., & Török, T. 2013, *The Astrophysical Journal*, 778, 99
- Leka, K. D. & Barnes, G. 2003a, *The Astrophysical Journal*, 595, 1277
- Leka, K. D. & Barnes, G. 2003b, *The Astrophysical Journal*, 595, 1296
- Leka, K. D. & Barnes, G. 2007, *The Astrophysical Journal*, 656, 1173
- Linton, M. G. & Antiochos, S. K. 2005, *The Astrophysical Journal*, 625, 506
- Liu, Y. & Schuck, P. W. 2012, *The Astrophysical Journal*, 761, 105
- Low, B. C. 1996, *Solar Physics*, 167, 217
- Lukin, V. S. & Linton, M. G. 2011, *Nonlinear Processes in Geophysics*, 18, 871
- Nindos, A. & Andrews, M. D. 2004, *The Astrophysical Journal*, 616, L175
- Pariat, E., Valori, G., Démoulin, P., & Dalmasse, K. 2015, *Astronomy and Astrophysics*, 580, A128
- Park, S.-H., Chae, J., & Wang, H. 2010, *The Astrophysical Journal*, 718, 43
- Park, S.-H., Cho, K. S., Bong, S.-C., et al. 2012, *The Astrophysical Journal*, 750, 48
- Park, S.-H., Lee, J., Choe, G. S., et al. 2008, *The Astrophysical Journal*, 686, 1397
- Phillips, A. D., MacNeice, P., & Antiochos, S. K. 2005, *The Astrophysical Journal*, 624, L129
- Pinto, R. F., Gordovskyy, M., Browning, P. K., & Vilmer, N. 2016, *Astronomy and Astrophysics*, 585, A159
- Prior, C. & Yeates, A. R. 2014, *The Astrophysical Journal*, 787, 100
- Rubio da Costa, F., Kleint, L., Petrosian, V., Liu, W., & Allred, J. C. 2016, *The Astrophysical Journal*, 827, 38
- Rust, D. M. 1994, *Geophysical Research Letters*, 21, 241
- Schrijver, C. J. 2007, *The Astrophysical Journal*, 655, L117
- Schrijver, C. J., De Rosa, M. L., Title, A. M., & Metcalf, T. R. 2005, *The Astrophysical Journal*, 628, 501
- Su, J. T., Jing, J., Wang, S., Wiegmann, T., & Wang, H. M. 2014, *The Astrophysical Journal*, 788, 150
- Taylor, J. B. 1974, *Physical Review Letters*, 33, 1139
- Tian, L. & Alexander, D. 2008, *The Astrophysical Journal*, 673, 532
- Tziotziou, K., Georgoulis, M. K., & Liu, Y. 2013, *The Astrophysical Journal*, 772, 115
- Tziotziou, K., Georgoulis, M. K., & Raouafi, N.-E. 2012, *The Astrophysical Journal Letters*, 759, L4
- Valori, G., Démoulin, P., & Pariat, E. 2012, *Solar Physics*, 278, 347
- Valori, G., Démoulin, P., Pariat, E., & Masson, S. 2013, *Astronomy and Astrophysics*, 553, 38
- Valori, G., Pariat, E., Anfinogentov, S., et al. 2016, *Space Science Reviews*, in press
- Vemareddy, P., Ambastha, A., Maurya, R. A., & Chae, J. 2012, *The Astrophysical Journal*, 761, 86
- Vemareddy, P. & Démoulin, P. 2016, arXiv.org, arXiv:1611.00699
- Woltjer, L. 1958, in *Proceedings of the National Academy of Sciences of the United States of America* (National Acad Sciences), 489–491
- Yang, S., Büchner, J., Santos, J. C., & Zhang, H. Q. 2013, *Solar Physics*, 283, 369
- Yang, S., Zhang, H. Q., & Büchner, J. 2009, *Astronomy and Astrophysics*, 502, 333
- Yang, X., Zhang, H. Q., Gao, Y., Guo, J., & Lin, G. 2012, *Solar Physics*, 280, 165
- Yu, D., Huang, X., Wang, H., et al. 2010, *The Astrophysical Journal*, 710, 869
- Zuccarello, F. P., Aulanier, G., & Gilchrist, S. A. 2015, *The Astrophysical Journal*, 814, 126

Appendix A: Gauge invariance

The helicity computation method of Valori et al. (2012) allows the estimation of helicity quantities using different sets of gauges. The potential vectors are computed using the DeVore gauge (DeVore 2000), i.e., $A_z = 0$. As discussed in Section 3.2 of Pariat et al. (2015), the derivation is performed by choosing a reference plane z_i , here either the top or bottom boundary of our dataset. At this boundary, Valori et al. (2012) proposed either choosing an integral or a derivative methodology. In the latter case, noted "DV-C" hereafter, \mathbf{A}_p , is simultaneously fulfilling the DeVore and Coulomb ($\nabla \cdot \mathbf{A}_p = 0$) gauge conditions. The former case will be simply noted "DV". Assuming $A_z(z_i) = A_{p,z}(z_i)$, we can therefore derive Eq. (7) using four different set of gauges, noted "DV bot", "DV top", "DV-C bot", and "DV-C top", with "top" and "bot" referring to the choice of z_i as respectively the top or bottom boundary.

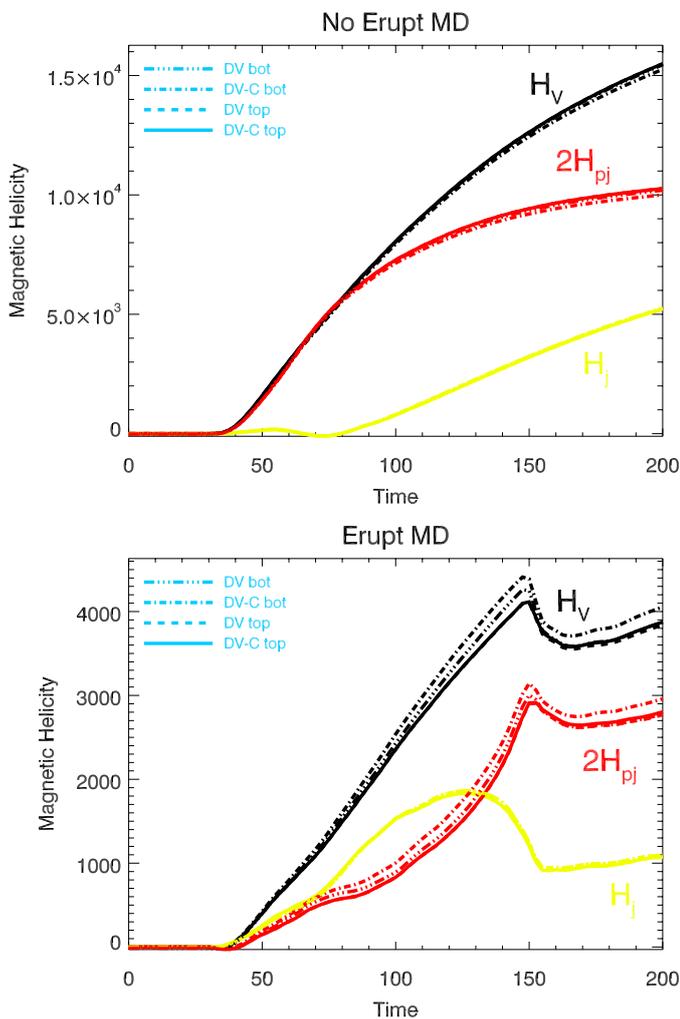


Fig. A.1. Helicity and its decomposition in the medium arcade simulations for the non-eruptive (top panel) and eruptive (bottom panel) cases. The [black, red, yellow] curves are respectively corresponding to $[H_V, 2H_{pj}, H_j]$. The helicities are computed using the resp. [DV bot, DV top, DV-C bot, DV-C top] gauge sets, corresponding to the resp. [three-dot-dashed, dashed, dot-dashed, continuous] curve.

Fig. A.1 displays the computation of H_V , H_j and H_{pj} in these four different gauges for both the eruptive and non-eruptive cases with the medium-strength arcade (MD) cases. One can remark that H_V , H_j and H_{pj} are gauge invariant. For the non-

eruptive case, the maximum relative standard deviation of the values of H_V is only 1% while it is a bit higher, 4%, for the eruptive simulations. The eruptive simulations indeed represent a more arduous case for helicity since both the bottom and the top boundary are significantly evolving in time. One observes that H_j is particularly gauge invariant with maximum standard deviation of 2% for both simulations. H_{pj} contains most of the gauge dependence of H_V with relative standard deviations which amount to 2% and 10% respectively for the non-eruptive and the eruptive cases. The other simulation cases present similar trends on the gauge dependence.

Overall, the distribution of values computed with different gauges provides us with rough estimates of the precision of the helicity computation. We can thus estimate that H_V has an average 5% measurement error. Any difference above this value can thus be considered as significant, which is the case of all the comparisons between simulations that are done in this paper. One can note that the gauges computed starting at the top boundary tend to be more consistent with each other, which is likely the results of the distribution of solenoidal errors specific to this simulation. As already noted in Valori et al. (2012) and Pariat et al. (2015), the DeVore-Coulomb approach is also less prone to numerical error. In the main text of this study, the helicity values are obtained using the "DV-C top" gauge set.

Appendix B: Magnetic helicity conservation

In addition to the volume helicity, the helicity computation method of Valori et al. (2012), allows for the derivation of the complete time integrated flux of helicity through, e.g., the 6 boundaries of the cartesian domain (cf. Pariat et al. 2015). This enables the determination of the helicity, $\Delta H_{\partial V}$, accumulating in the domain in time:

$$\Delta H_{\partial V}(t) = \int_0^t F_{\text{tot}}(\tau) d\tau \quad (\text{B.1})$$

with F_{tot} the flux of helicity given by equation (30) of Pariat et al. (2015). The computation of $\Delta H_{\partial V}$ requires both the knowledge of the velocity and magnetic field distribution on ∂V . It is independent of the derivation of H_V . $\Delta H_{\partial V}$ is theoretically gauge invariant (Pariat et al. 2015). Fig. B.1 presents the measurements of $\Delta H_{\partial V}$ with the four gauge sets considered in Appendix A, for the medium arcade cases. Our results shows indeed that $\Delta H_{\partial V}$ is relatively gauge invariant, with a 1% maximum relative standard deviation of the values for the non-eruptive simulation. Here again the eruptive case shows slightly larger error, with a 6% relative standard deviation. As already noted in Pariat et al. (2015), because the computation of $\Delta H_{\partial V}$ requires more integration, and in particular a time integration, it tends to be more sensitive to numerical errors than H_V and therefore presents a slightly stronger gauge dependence. The gauge sets computed from the top boundary are also more consistent, being less sensitive to errors

In ideal MHD, magnetic helicity is a conserved quantity (Woltjer 1958). Even when non-ideal process, such as magnetic reconnection, are developing, magnetic helicity is believed to be a quasi-conserved quantity (Taylor 1974; Berger 1984). In a recent numerical simulation of a solar-like active event, with intense magnetic reconnection, Pariat et al. (2015) have shown that magnetic helicity was extremely well conserved, with a $< 2\%$ dissipation, particularly with regards to the magnetic energy dissipation which was more than 30 times higher. Following Yang et al. (2013) and Pariat et al. (2015), the comparison of the time-integrated helicity flux and volume helicity enables one study the

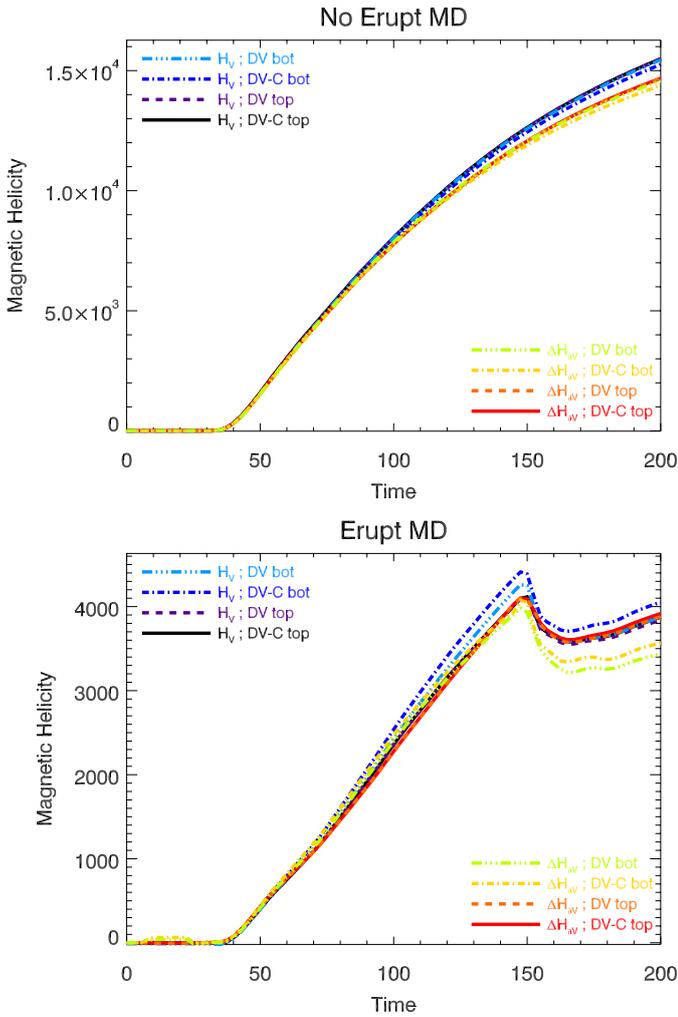


Fig. B.1. Helicity variation and accumulated flux in the medium arcade simulations for the non-eruptive (left panel) and eruptive (right panel) cases. The volume helicity estimations, H_V , are plotted in black, purple, blue and cyan while the time-integrated flux of helicity, $\Delta H_{\delta V}$, are plotted in the red, orange, yellow and green. The helicities are computed using the resp. [DV bot, DV top, DV-C bot, DV-C top] gauge sets, corresponding to the resp. [three-dot-dashed, dashed, dot-dashed, continuous] curve.

helicity conservation in the numerical simulation, and how close the numerical scheme is to ideal MHD behavior. In the present simulations, since the helicity in the system is initially null, assuming perfect helicity conservation, $\Delta H_{\delta V}$ and H_V should be equal. For the medium arcade cases, one observes on Fig. B.1 that, independently of the gauge set used, the curves of $\Delta H_{\delta V}$ and of H_V match well. This indicates a good level of conservation of magnetic helicity in these simulations. Using the relative accumulated helicity difference criteria, ϵ_H , defined in equation (43) of Pariat et al. (2015), we note maximum values, using the least-conservative gauge set, of 6% for the non-eruptive simulation and 15% for the eruptive one. This demonstrates that while not perfect, the code is able to conserve relatively well magnetic helicity. Similar helicity-conservation trends are obtained for all the simulation datasets used in this study.